



LAWRENCE
LIVERMORE
NATIONAL
LABORATORY

The galaxy hosts and large-scale environments of short-hard (gamma)-ray bursts

J. X. Prochaska, J. S. Bloom, H.-W. Chen, R. J. Foley, D. A. Perley, E. Ramirez-Ruiz, J. Granot, W. H. Lee, D. Pooley, K. Alatalo, K. Hurley, M. C. Cooper, A. K. Dupree, B. F. Gerke, B. M. S. Hansen, J. S. Kalirai, J. A. Newman, R. M. Rich, H. Richer, S. A. Stanford, D. Stern, W. J. M. van Breugel

April 12, 2006

The Astrophysical Journal

Disclaimer

This document was prepared as an account of work sponsored by an agency of the United States Government. Neither the United States Government nor the University of California nor any of their employees, makes any warranty, express or implied, or assumes any legal liability or responsibility for the accuracy, completeness, or usefulness of any information, apparatus, product, or process disclosed, or represents that its use would not infringe privately owned rights. Reference herein to any specific commercial product, process, or service by trade name, trademark, manufacturer, or otherwise, does not necessarily constitute or imply its endorsement, recommendation, or favoring by the United States Government or the University of California. The views and opinions of authors expressed herein do not necessarily state or reflect those of the United States Government or the University of California, and shall not be used for advertising or product endorsement purposes.

The galaxy hosts and large-scale environments of short-hard γ -ray bursts

J. X. Prochaska¹, J. S. Bloom², H.-W. Chen³, R. J. Foley², D. A. Perley², E. Ramirez-Ruiz^{4,1}, J. Granot^{5,4}, W. H. Lee⁶, D. Pooley^{2,7}, K. Alatalo², K. Hurley⁸, M. C. Cooper², A. K. Dupree⁹, B. F. Gerke¹⁰, B. M. S. Hansen¹¹, J. S. Kalirai¹, J. A. Newman^{12,7}, R. M. Rich¹¹, H. Richer¹³, S. A. Stanford^{14,15}, D. Stern¹⁶, W.J.M. van Breugel¹⁵

¹*UCO/Lick Observatory, University of California, Santa Cruz; Santa Cruz, CA 95064*

²*Department of Astronomy, 601 Campbell Hall, University of California, Berkeley, CA 94720-3411.*

³*Center for Space Research, Massachusetts Institute of Technology, Cambridge, MA 02139-4307*

⁴*Institute for Advanced Study, Olden Lane, Princeton, NJ 08540*

⁵*KIPAC, Stanford University, P.O. Box 20450, Mail Stop 29, Stanford, CA 94309*

⁶*Instituto de Astronomia, UNAM Apdo. Postal 70-264 Mexico DF 04510 Mexico*

⁷*Hubble Fellow*

⁸*UC Berkeley, Space Sciences Laboratory, 7 Gauss Way, Berkeley, CA 94720-7450*

⁹*Harvard-Smithsonian Center for Astrophysics, 60 Garden St., Cambridge, MA 02138*

¹⁰*Department of Physics, 366 LeConte Hall, University of California, Berkeley, CA 94720-7300*

¹¹*Department of Physics and Astronomy, University of California Los Angeles, Los Angeles, CA, 90095*

¹²*Institute for Nuclear and Particle Astrophysics, Lawrence Berkeley National Laboratory, Berkeley, CA 94720*

¹³*Physics & Astronomy Department, University of British Columbia, Vancouver, B.C. V6T 1Z1*

¹⁴*Department of Physics, One Shields Ave., University of California, Davis, CA 95616-8677*

¹⁵*Institute for Geophysics and Planetary Physics, Lawrence Livermore National Laboratory, L-413 7000 East Ave, Livermore CA 94550*

¹⁶*Jet Propulsion Laboratory, California Institute of Technology, 4800 Oak Grove Drive, MS 169-506, Pasadena, CA*

The nature of the progenitors of short duration, hard spectrum, gamma-ray bursts¹ (GRBs) has remained a mystery. Even with the recent localizations of four short-hard GRBs, no transient emission has been found at long wavelengths that directly constrains the progenitor nature. Instead, as was the case in studying the different morphological subclasses of supernovae^{2,3} and the progenitors of long-duration GRBs⁴, we suggest that the progenitors of short bursts can be meaningfully constrained by the environment in which the bursts occur. Here we present the discovery spectra of the galaxies that hosted three short-hard GRBs and the spectrum of a fourth host. The results indicate that these environments, both at the galaxy scale and galaxy-cluster scale, differ substantially from those of long-soft GRBs. The spatial offset of three bursts from old and massive galaxy hosts strongly favours an origin from the merger of compact stellar remnants, such as double neutron stars or a neutron-star black hole binary. The star-forming host of another GRB provides confirmation that, like supernovae of Type Ia, the progenitors of short-hard bursts are created in all galaxy types. This indicates a class of progenitors with a wide distribution of delay times between formation and explosion.

In the past four months the Swift and HETE-II satellites have discovered four GRBs whose short duration ($t < 2$ s) and spectral hardness place them within the short-hard GRB classification^{5–8}. Furthermore, each of these GRBs has been localized by its afterglow X-ray emission to within a circle of radius $10''$ on the sky^{9–12}. Although previous missions reported hundreds of short-hard GRBs, none of these were promptly localized to less than a few arcminutes and so a counterpart

association at other wavelengths proved elusive^{13,14}. The discovery of GRB 050509b and a fading X-ray afterglow⁵ led to the first redshift and host galaxy association⁹ for a short-hard GRB, solving the long-standing mystery over the distance scale and energetics for at least some members of this class. The four events now localized offer an opportunity to study the population of host galaxies and large-scale environments, examine the energetics, and begin to constrain the nature of the progenitors.

Based on positions of the afterglows, two of four bursts (050509b and 050813) are associated with clusters of galaxies^{9,15}. Because only $\approx 10\%$ of the mass of the Universe is contained within clusters, this suggests that either galaxies in clusters preferentially produce progenitors of short-hard GRBs or that short-hard bursts are preferentially more likely to be discovered and localized in cluster environments⁹. We have examined the Swift X-ray Telescope data of the fields of the other two GRBs (050709 and 050724) and found no conclusive evidence for diffuse hot gas associated with massive clusters. Furthermore, a spectroscopic study of three bright galaxies near the X-ray afterglow position of GRB 050724 show them all to be at different redshifts, disfavours a cluster origin for that burst. The cluster environments of at least two short-hard GRBs contrast strikingly with the observation that no well-localized long-soft GRB has yet been associated with a cluster¹⁶. Therefore, more sensitive observations of the fields of both historical and new well-localized short-hard GRBs may be expected to show a significant preponderance to correlate with galaxy clusters.

We now turn to the putative galaxy hosts of short-hard GRBs. In three of four cases, the GRB has been plausibly associated with a galaxy to better than a 99% confidence level (Figure 1). In the

fourth case (050813), there are two galaxies located in the error circle with comparable magnitude and one may associate the event with either of these. In Figure 2, we present the discovery spectra of three short-hard host galaxies and a high-resolution spectrum of GRB 050709 which was previously identified¹⁷. Three bursts are associated with galaxies exhibiting characteristic ‘early-type’ spectra. The absence of observable $H\alpha$ and $[O\ II]$ emission constrains the unobscured star formation rates (SFR) in these galaxies to $SFR < 0.2 M_{\odot} \text{yr}^{-1}$ (Table 1), where M_{\odot} is the mass of the Sun, and the lack of Balmer absorption lines implies that the last significant star forming event occurred > 1 billion years ago. The host galaxy of GRB 050709 exhibits strong emission lines that indicate on-going star formation with a conservative lower limit of $SFR > 0.5 M_{\odot} \text{yr}^{-1}$. These observations clearly indicate that these short-hard GRBs occurred during the past ~ 7 billion years of the Universe in galaxies with diverse physical characteristics.

In contrast to what is found for short-hard GRBs, all of the confirmed long-soft GRB host galaxies are actively forming stars with integrated, unobscured SFRs $\approx 1 - 10 M_{\odot} \text{yr}^{-1}$ ¹⁸. The galaxies have small stellar masses and bluer colors than present-day spiral galaxies¹⁹ (suggesting a low metallicity). We therefore conclude that the host galaxies of short-hard GRBs, and by extension the progenitors, are not drawn from the same parent population of long-soft GRBs. And although long-soft GRBs are observed to significantly higher redshift than the current short-hard GRB sample, one reaches the same conclusions when restricting to low- z long-soft GRB hosts²⁰.

The identification of three galaxies without current star formation argues that the accepted progenitor model of long-soft GRBs (the collapse of a massive star²¹) is unviable as a source for

the short-hard GRBs. Instead, the observations lend support to theories in which the progenitors of short-hard GRBs are merging compact binaries (neutron stars or black holes^{22,23}). This inference is supported through several channels. First, the redshift distribution of these short-hard bursts is inconsistent with a bursting rate that traces the star-formation rate in the universe, unlike long-soft GRBs, which do follow it. If we introduce a ~ 1 Gyr time delay from starburst to explosion, as expected from compact object mergers, the observed redshift distribution of these GRBs is consistent with the star-formation rate²⁴. Second, the lack of an associated supernova for all four short-hard GRBs is strong evidence against a core-collapse origin. Third, our measured offsets (fig. 1) of the short-hard GRBs from their putative hosts are compatible with predicted site of merging compact remnant progenitors^{25,26}. Noteworthy, and somewhat counterintuitive, is that the albeit small offset of GRB 050724 (2.36 ± 0.90 kpc) is near the median predicted merger offset for such galaxies²⁶.

The identification of the host galaxies and redshifts finally fixes the absolute burst energies. Table 2 shows the inferred isotropic energy release in prompt γ -ray emission, along with its duration in the source rest-frames. These events suggest that short-hard GRBs are less energetic, typically by more than one order of magnitude, than their long counterparts, which typically release a total γ -ray energy of 5×10^{50} erg when collimation is taken into account. The total isotropic-equivalent energy in γ -rays, $E_{\gamma,\text{iso}}$ appears to correlate with the burst duration, such that longer events are also more powerful²⁷. We find that $E_{\gamma,\text{iso}} \propto T_{90}^{\psi}$ and $\psi \approx 3/2$ to 2. The total energies, durations, and the general behavior of the correlation between them are in rough agreement with the numerical modeling of GRB central engines arising from compact object mergers²⁸. Our fits

to the available afterglow data indicate that the density in the circumburst medium is closer to that found in the interstellar ($n \approx 1 \text{ cm}^{-3}$) rather than intergalactic medium ($n \approx 10^{-3} \text{ cm}^{-3}$). This might suggest a selection bias where short-hard GRBs that occur in a dense external medium have a brighter afterglow emission, and thus are more accurately localized⁹.

The association of short-hard GRBs with both star-forming galaxies and with ellipticals dominated by old stellar populations is analogous to type Ia SNe. It indicates a class of progenitors with a wide distribution of delay times between formation and explosion, with a tail probably extending to many Gyr. Similarly, just as core-collapse supernovae are discovered almost exclusively in late-time star-forming galaxies, so too are long-soft GRBs. As new redshifts, offsets and host galaxies of short-hard GRBs are gathered, the theories of the progenitors will undoubtedly be honed. Still, owing to the largely featureless light of afterglow radiation, unless short-hard bursts are eventually found to be accompanied by tell-tale emission features like the supernovae of long-duration GRBs, the only definitive understanding of the progenitors will come with the observations of concurrent gravitational radiation or neutrino signals arising from the dense, opaque central engine.

1. Kouveliotou, C. *et al.* Identification of two classes of gamma-ray bursts. *ApJ (Letters)* **413**, 101–104 (1993).
2. Reaves, G. Notes on extragalactic supernovae. *PASP* **65**, 242 (1953).
3. van Dyk, S. D. Association of supernovae with recent star formation regions in late type galaxies. *AJ* **103**, 1788–1803 (1992).

4. Bloom, J. S., Kulkarni, S. R. & Djorgovski, S. G. The Observed Offset Distribution of Gamma-Ray Bursts from Their Host Galaxies: A Robust Clue to the Nature of the Progenitors. *AJ* **123**, 1111–1148 (2002).
5. Gehrels, N. *et al.* The first localization of a short gamma-ray burst by Swift (2005). [astro-ph/0505630](#).
6. Butler, N. *et al.* GRB050709: a possible short-hard GRB localized by HETE. (2005). GCN Circular 3570.
7. Covino, S. *et al.* GRB050724: a short-burst detected by swift. (2005). GCN Circular 3665.
8. Sato, G. *et al.* GRB050813: Swift-BAT refined analysis (2005). GCN Circular 3793.
9. Bloom, J. S. *et al.* Closing in on a Short-Hard Burst Progenitor: Constraints from Early-Time Optical Imaging and Spectroscopy of a Possible Host Galaxy of GRB 050509b (2005). [astro-ph/0505480](#).
10. Fox, D. B., Frail, D. A., Cameron, P. B. & Cenko, S. B. GRB 050709: Candidate X-ray Afterglow from Chandra (2005). GCN Circular 3585.
11. Burrows, D. N. *et al.* GRB 050724: Chandra Observations of the X-ray Afterglow (2005). GCN Circular 3697.
12. Morris, D. C. *et al.* GRB 050813: Swift XRT afterglow localization (2005). GCN Circular 3790.

13. Hurley, K. *et al.* Afterglow Upper Limits for Four Short-Duration, Hard Spectrum Gamma-Ray Bursts. *ApJ* **567**, 447–453 (2002).
14. Nakar, E., Gal-Yam, A., Piran, T. & Fox, D. B. The distances of short-hard GRBs and the SGR connection (2005). astro-ph/0502148.
15. Gladders, M., Berger, E., Morrell, N. & Roth, M. Grb 050813: Magellan detection of a high redshift galaxy cluster (2005). GCN Circular 3798.
16. Bornancini, C. G. *et al.* The Galaxy Density Environment of Gamma-Ray Burst Host Galaxies. *ApJ* **614**, 84–90 (2004).
17. Price, P. A., Roth, K. & Fox, D. W. GRB 050709: spectroscopy. (2005). GCN Circular 3605.
18. Christensen, L., Hjorth, J. & Gorosabel, J. UV star-formation rates of GRB host galaxies. *A&A* **425**, 913–926 (2004).
19. Le Floc’h, E. *et al.* Are the hosts of gamma-ray bursts sub-luminous and blue galaxies? *A&A* **400**, 499–510 (2003).
20. Sollerman, J. *et al.* On the nature of nearby GRB/SN host galaxies (2005). astro-ph/0506686.
21. Woosley, S. E. Gamma-ray bursts from stellar mass accretion disks around black holes. *ApJ* **405**, 273–277 (1993).
22. Paczynski, B. Gamma-ray bursters at cosmological distances. *ApJ (Letters)* **308**, L43–L46 (1986).

23. Eichler, D., Livio, M., Piran, T. & Schramm, D. N. Nucleosynthesis, neutrino bursts and gamma-rays from coalescing neutron stars. *Nature* **340**, 126–128 (1989).
24. Guetta, D. & Piran, T. The luminosity and redshift distributions of short-duration GRBs. *A&A* **435**, 421–426 (2005).
25. Fryer, C. L., Woosley, S. E. & Hartmann, D. H. Formation Rates of Black Hole Accretion Disk Gamma-Ray Bursts. *ApJ* **526**, 152–177 (1999).
26. Bloom, J. S., Sigurdsson, S. & Pols, O. R. The spatial distribution of coalescing neutron star binaries: implications for gamma-ray bursts. *MNRAS* **305**, 763–769 (1999).
27. Berger, E. *et al.* A merger origin for short gamma-ray bursts inferred from the afterglow and host galaxy of grb 050724 (2005). astro-ph/0508115.
28. Lee, W. H., Ramirez-Ruiz, E. & Granot, J. A compact binary merger model for the short, hard GRB 050509b (2005). astro-ph/0506104.
29. Eisenstein, D. J., Hogg, D. W. & Padmanabhan, N. GRB050509b, SDSS pre-burst observations. *GRB Circular Network* **3418**, 1–+ (2005).
30. Berger, E. GRB 050813: Gemini spectroscopy and redshifts of galaxies B and C (2005). GCN Circular 3801.

Acknowledgements We thank S. Sigurdsson and D. Kocevski for useful discussions. Some of these observations were made with the W.M. Keck Telescope. The Keck Observatory is a joint facility of the University of California, the California Institute of Technology, and NASA. We are grateful to the staff of

Gemini for their assistance in acquiring this data. J.X.P., J.S.B., and H.-W.C. are partially supported by NASA/Swift grant NNG05GF55G. Work at LLNL is performed under the auspices of the U.S. Department of Energy by UC, Lawrence Livermore National Laboratory under contract No. W-7405-Eng-48.

Competing Interests The authors declare that they have no competing financial interests.

Correspondence Correspondence and requests for materials should be addressed to J.X.P.

(email: xavier@ucolick.org).

Physical Properties of the Hosts of Short-Hard GRBs

GRB	z	r^a (kpc)	R^b (mag)	L_B^c ($10^9 L_\odot$)	SFR ^d ($M_\odot \text{yr}^{-1}$)	Spectral Type
050509b	0.2248 ± 0.0002	39 ± 13	16.8 ± 0.05	20	< 0.1	Elliptical
050709	0.1606 ± 0.0001	3.5 ± 1.3	21.1 ± 0.2	0.4	> 0.5	Late-type dwarf
050724	0.2576 ± 0.0004	2.4 ± 0.9	19.8 ± 0.3	2	< 0.05	Early-type
050813 (B)	0.719 ± 0.001	...	23.43 ± 0.07	4	< 0.1	Elliptical
050813 (C)	0.73 ± 0.01	...	22.57 ± 0.07	10	< 0.2	Elliptical
050813 (X)	0.722 ± 0.001	...	22.75 ± 0.07	8	< 0.1	Elliptical

^aProjected offset of the X-ray afterglow positions from the optical centroid of the respective host galaxies. The quoted error is an approximation to the uncertainty of the most likely offset r , following appendix B of ⁴, which is required because offsets are a positive-definite quantity and not strictly Gaussian. In general, $r \pm \sigma_r$ does not contain 68% of the probability distribution function.

^b R -band magnitudes. We convert the Sloan Digital Sky Survey r magnitude for 050509b²⁹. For the galaxies associated with GRB 050813 we have measured i -band magnitudes and converted to R -band assuming $R - i = 0.99$ mag, appropriate for an elliptical galaxy at $z = 0.7$.

^cThe R -band magnitudes were converted to B -band luminosities by assuming standard colors for these spectral types, adopting the redshift listed in column 1, and adopting the standard cosmology $\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$ and Hubble's constant $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$. The luminosities have not been corrected for Galactic extinction and are reported relative to the Solar B -band

luminosity.

^dUnextincted star formation rate based on $H\alpha$ and/or [OII] luminosity. Upper limits are 3σ .

Inferred Burst Energetics and Durations

GRB	$E_{\gamma,\text{iso}}[\text{erg}]^a$	$T_{90}/(1+z) [\text{sec}]^b$
050509b	2.75×10^{48}	0.032
050709	2.29×10^{49}	0.060
050724	1.0×10^{50}	0.203
050813	1.7×10^{50}	0.349

^a Isotropic-equivalent energy $E_{\gamma,\text{iso}}$, computed using the observed fluence and redshift under the assumption of a concordance cosmology with $\Omega_m = 0.29$, $\Omega_\Lambda = 0.71$ and Hubble's constant $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$. While these energies are systematically lower than for long-soft GRBs, we note that with the energy range covered by Swift (15–350 keV) and the spectral properties of the prompt emission, the derived values should be considered lower limits.

^b Source rest-frame duration, measured in T_{90} , the time when 90% of the total fluence of the GRB is accumulated, beginning after 5% of the fluence has been accumulated¹.

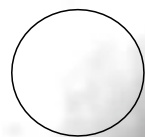
Figure 1 Optical light montage of four host galaxy regions of short-hard GRBs. In the case of GRB 050709 and GRB 050724 where optical afterglows were detected, the GRB is projected to within $2''$ from the center of a galaxy with apparent magnitude $R < 19.5$ mag. The likelihood of a chance association between these afterglows and the putative host galaxies is less than 10^{-4} per event given the covering fraction of such objects on the sky. Similarly, the error circle containing GRB 050509b encompasses a single bright galaxy which is the putative host galaxy ⁹ for which the chance of a spurious physical association with the burst is $\sim 10^{-3}$. Images were acquired on the Gemini North 8m Telescope (GRB 050724, i-band; GRB 050813) and Keck 10m Telescopes (GRB 050509b, R-band; GRB 050709, R-band) and processed in the usual manner. Processed images were registered to an absolute world coordinate system with typical 1σ rms uncertainties of 150 milliarcsecond in each coordinate. We find the absolute positions of host galaxies for 050509b, 050709, and 050724 as $\alpha(\text{J2000}) = 12:36:12.878$ $\delta(\text{J2000}) = +28:58:58.95$, $23:01:26.849 -38:58:39.39$, and $16:24:44.381 -27:32:26.97$, respectively. The ellipses in each panel represent the astrometric position of the most accurate X-ray afterglow position reported (68% confidence interval for GRB 050509b⁹; 68% confidence interval for GRB 050709 ¹⁰; 68% confidence interval for GRB 050724 ¹¹; and reflect the uncertainty in the astrometric tie between the X-ray and optical frame. The 90% containment radius previously reported for GRB 050813¹² is shown as a large circle. With the same data, using an optimized technique for faint transient localization⁹, we have localized GRB 050813 to $\alpha(\text{J2000}) = 16:07:56.953 \pm 0.20$ sec, $\delta(\text{J2000}) = +11:14:56.60 \pm 1.45$ arcsec. The

smaller ellipse shows this 68% containment radius. This localization makes the host identification of B or even the fainter B^* more likely over galaxy C . Adopting the redshift of the putative host or cluster redshift (GRB 050813) a projection scale is shown at right in each panel. The galaxies labeled in the panel GRB 050813 panel are referred to in figure 2. We note that galaxies X (16:07:57.509 +11:15:02.13; $i = 21.76 \pm 0.03$ mag), B (16:07:57.200 +11:14:53.09; $i = 22.44 \pm 0.04$), and C (16:07:57.008 +11:14:47.37; $i = 21.58 \pm 0.04$) show consistent, red colors that suggest a cluster membership¹⁵. The brightest objects at the edge of the large error circle (16:07:57.393 +11:14:42.79 and 16:07:56.850 +11:15:01.12) are likely foreground Galactic stars. All images were smoothed with a Gaussian of 1.4–1.6 pixels to enhance the contrast between detected objects and sky noise. North is up and East is to the left.

Figure 2 Optical spectroscopy for the host galaxies of short-hard GRBs. With the exception of GRB 050724, these data are the discovery spectra which established the redshift of the GRB event and also the properties of the galaxy host and/or environment. The data were acquired with the (a) Echellette Spectrometer and Imager on Keck II with a 1'' slit in echellette mode; (b) the DEIMOS spectrometer on Keck II obtained through a 0.7'' longslit using the 600line/mm grating; (c) the LRIS spectrometer on the Keck I telescope with the 600/4000 grism through a 1'' longslit for the blue spectrum and the GMOS spectrometer on the Gemini-North telescope using a 0.75'' slit (following astrometry based on a Magellan guide-camera image) and the R400 grating centered at 690nm for the red spectrum; and (d) the GMOS spectrometer using the identical setup as (c). The data were

fluxed using spectrophotometric standards taken with the same instrumental setups. The absolute flux is uncertain, in particular, due to slit losses and is not corrected for reddening by the Milky Way. The redshifts of the galaxies were measured through fits to the spectral features indicated in the figure. We obtained spectra of two bright galaxies near GRB 050724 (at positions 16:24:46.739 -27:32:28.90 and 16:24:43.344 -27:32:07.21) and did not find them to be at the same redshift as the host galaxy; we therefore have found no evidence the GRB 050724 is a member of a galaxy cluster. Note that we present only the spectrum for galaxy B associated with GRB 050813 (figure 1). Our spectrum of galaxy C shows a 4000Å break consistent with $z = 0.73$ and no significant emission lines, galaxy X shows absorption features indicating $z = 0.722$ (see also ³⁰), and we have no redshift constraint for galaxy B^* ($i = 24.2 \pm 0.1$). The small projected distance between these sources ($\approx 40 - 100 h_{70}^{-1}$ kpc) and large velocity difference ($\Delta v = 690 - 3000$ km s⁻¹) strongly support the cluster nature of the progenitor environment for GRB050813¹⁵.

GRB 050509b



30 kpc

$z=0.2248$

GRB 050709



30 kpc

$z=0.1606$

GRB 050724

30 kpc

$z=0.2576$

[$z=0.719$]

x

B

C

60 kpc

GRB 050813

